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Thermal expansion of liquid Fe-S alloy at high pressure

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Highlights

- Density of liquid Fe-S alloys has been measured under high pressure as a function of temperature by *in situ* X-ray diffraction in multi-anvil press.
- Thermal expansion of liquid Fe-S alloys has been determined up to 7 GPa and 2200 K.
- Top-down crystallization is the most likely scenario for Fe-FeS cores of planetesimals and small planets.

Abstract

Local structure and density of liquid Fe-S alloys at high pressure have been determined *in situ* by combined angle and energy dispersive X-ray diffraction experiments in a multi-anvil apparatus, covering a large temperature and compositional range. Precise density measurements collected for increasing temperature allowed us to directly derive the thermal expansion coefficients for liquid Fe-S alloys as a function of composition. In turn, thermal expansion has been used to refine thermodynamic models and to address the crystallization regime of telluric planetary cores by comparing the adiabatic temperature gradient and the slope of the liquidus in the Fe-FeS system.

For Fe-S cores of asteroids and small planetesimals, top-down solidification is the dominant scenario as the compositional domain for which the slope of the liquidus is greater than the adiabatic gradient is limited to a narrow portion on the Fe-rich side. However, bottom-up growth of the inner core is expected for S-poor cases, with this compositional domain expanding to more S-rich compositions with increasing pressure (size of the planetary body). In particular, bottom-up crystallization cannot be excluded for the Moon and Ganymede.

1. Introduction

Many recent studies highlight a rich diversity in the core structure of terrestrial bodies, largely depending on the core crystallization regime (e.g. Breuer et al., 2015; Dumberry and Rivoldini, 2015; Hauck et al., 2006; Rivoldini et al., 2011; Rückriemen et al., 2018, 2015; Williams, 2009). This, in turn, directly depends on pressure conditions and light element content, which relates to accreting material, planetary differentiation process and planet's size. From an observational point of view, indications on the core structure and dynamics can be derived from geodetic data (static and dynamic gravity field, and rotation) and, when present, from the magnetic field. The presence of an internally generated magnetic field is relatively common in the terrestrial bodies of our solar system. Aside from Earth, active global magnetic fields of internal origin have been detected on Mercury and Ganymede (Kivelson et al., 1997; Ness, 1979). A strong, though now extinct, magnetic field was also present on Mars, as suggested by crustal magnetic field maps (e.g. Acuña et al., 1999, 2001; Connerney et al., 2001) and on the Moon, as evidenced from the remnant magnetism of the lunar crust (e.g. Weiss and Tikoo, 2014). Furthermore, magnetized meteorites indicate that planetesimals and numerous planetary bodies may have had their own self-generated, long-lived fields (e.g. Weiss et al., 2008; Shah et al., 2017; Bryson et al., 2019a). A dynamo operating in the fluid metallic core is the most likely mechanism for generating a planetary magnetic field, and compositional convection driven by core solidification is one of the main power sources for a long standing dynamo (Stevenson, 2003). The strength and lifetime of a magnetic field is thus directly tied to the crystallization process of the core, which can proceed very differently in various planets, depending on composition and size (e.g. Breuer et al., 2015; Rückriemen et al., 2018).

As a first order guidance, neglecting complexities arising because of nucleation barriers (e.g. Davies et al., 2019; Huguet et al., 2018), we here assume that solid inner core starts to form when the temperature drops below the melting curve of the core-forming material (the liquidus). In a convecting medium the temperature profile is almost isentropic and the core temperature is generally assumed to follow an adiabat. The relative slopes (dT/dP) of the adiabat and of the liquidus thus determine the style of core solidification (e.g. Breuer et al., 2015; Williams, 2009). If the adiabat is shallower than the liquidus, crystallization occurs at depth, as for the Earth, and the inner core grows bottom-up. On the contrary, if the adiabat is steeper than the liquidus, top-down solidification is expected, in which shallow crystallization occurs and solid material sinks or floats, depending upon density contrast with the liquid, leading to possible compositional stratification (e.g. Dumberry and Rivoldini, 2015; Hauck et al., 2006; Rückriemen et al., 2018, 2015). Although both scenarios don't rule out the development of a dynamo, how and when a magnetic field can be generated in the course of top-down crystallization remains debated (e.g. Breuer et al., 2015).

Sulfur (S) is classically considered to be the major light element alloyed to iron (Fe) in the core of small planetary bodies (e.g. Antonangeli et al., 2015; Breuer et al., 2007; Morard et al., 2018; Terasaki et al., 2019; Rückriemen et al., 2015; Weber et al., 2011). The actual S content depends on primordial composition of the planetary forming material and the partitioning of sulfur between silicate and metallic phases, which in turn depends on

temperature and pressure conditions, as well as on the oxidation state, during core formation. For instance, while both Mercury and Mars are thought to have significant S in their bulk composition, the greater distance from the Sun and the more oxidizing formation conditions support more significant enrichment of S in the core of Mars than in that of Mercury (Dumberry and Rivoldini, 2015; Hauck et al., 2006; Rivoldini et al., 2011). Ultimately differences in the P, T, fO_2 conditions of core formation result in a large variation in the S abundance in the liquid cores of planets and moons. For instance, S content in the Moon's core is classically limited to less than 8 wt.% (e.g. Antonangeli et al., 2015; Laneuville et al., 2014; Rai and Van Westrenen, 2014; Steenstra et al., 2017), while in the case of the Jovian moons Ganymede and Callisto, the core might have a S concentration greater than the eutectic concentration (e.g. Scott et al., 2002). As a direct consequence of this composition diversity, different crystallization scenarios can arise. Indeed S content significantly affects the difference in the adiabat vs. liquidus relations and density contrast between the crystallizing solid and the liquid phases (e.g. Breuer et al., 2015; Williams, 2009).

Thermal expansion is a central parameter controlling the slope of the adiabat (Stacey, 2005; Williams, 2009), which can be written as

$$dT/dP = \alpha(P)T / \rho(P)C_P \quad (1)$$

where $\alpha(P)$ is the pressure-dependent thermal expansion, T the temperature at which the adiabatic gradient is calculated, $\rho(P)$ the pressure-dependent density, and C_P the heat capacity at constant pressure.

According to its definition, thermal expansion can be experimentally derived from the variation of the density with the temperature (at constant pressure)

$$\alpha(P,T) = -1/\rho(P,T) (d\rho(P)/dT) \quad (2)$$

Due to the experimental difficulties in density measurement, neither thermal expansion measurements over a wide compositional range (0-50 at% S) nor under high pressure have been performed. Thermal expansion of liquids in the Fe-FeS system has been studied exclusively at ambient pressure and only for end-member liquids, Fe and FeS (e.g. Assael et al., 2006; Kaiura and Toguri, 1979; Nagamori, 1969). Results on FeS (Kaiura and Toguri, 1979) are limited to few points covering a very limited T range (~1500-1650 K). Even in the case of liquid Fe, in spite of the large number of studies, proposed values for thermal expansion show large discrepancies (e.g.; 11.0×10^{-5} K^{-1}, Hixson et al., 1990; 8.2×10^{-5} K^{-1}, Nasch and Steinemann, 1995; 13.2×10^{-5} K^{-1}, Assael et al., 2006). Density determinations of liquid Fe-S alloys under high pressure includes ex-situ sink-float (Balog et al., 2003; Nishida et al., 2008), X-ray absorption (Chen et al., 2014; Nishida et al., 2011; Sanloup et al., 2000; Terasaki et al., 2019) and X-ray diffraction (Morard et al., 2018) methods. However, none of these methods provided sufficient temperature-dependent density data to allow precise determination of thermal expansion, due to combined difficulties in density measurements and temperature control at high pressure.

As highlighted by Williams (2009), extrapolation of thermal expansion of liquid Fe and FeS to pressures, temperatures and compositions directly relevant for planetary cores comes

with significant uncertainties that limit the reliability of the assessment of the relative slopes of the adiabats and liquidus curves (which also is not so well known). The determination of thermal expansion of liquid Fe-S alloys under the high-pressure conditions existing in planetary cores is thus of fundamental importance for refining thermodynamic models of planetary cores (Morard et al., 2018; Rivoldini et al., 2011; Terasaki et al., 2019) and to constrain adiabatic heat flux in cores of telluric bodies (Silber et al., 2018).

In this study, we provide direct determinations of the thermal expansion of liquid Fe-S alloys in the 4 to 9 GPa range, pressures directly relevant for the core of the Moon and Ganymede, using a combined angular and energy dispersive X-ray diffraction (CAESAR) technique (Wang et al., 2004). X-ray diffuse scattering from a liquid sample was analyzed to extract local structure information and also used to determine liquid density at high pressure (Eggert et al., 2002; Morard et al., 2014). To date, this method has been applied in Paris-Edinburgh (PE) cell and diamond anvil cell (DAC) for structural and density studies of liquids under high pressure (e.g. Morard et al., 2018, 2017, 2014; Sanloup et al., 2013; Yamada et al., 2011, 2007). Here we present the nontrivial extension of these measurements to the multi-anvil press. Noticeably, extension of this methodology to the multi-anvil press enables investigations over larger pressure range than achievable in PE experiments and, more importantly, the ability to vary the temperature of a confined liquid with the precision needed for the determination of thermal expansion, refinements not possible in the above-mentioned early PE and DAC studies. Measured thermal expansions are used together with thermodynamic modeling to assess the pressure and compositional dependence of the adiabatic gradients in the liquid Fe-FeS system in the 0 to 10 GPa range and, by comparison with the slope of the liquidus, to discuss the crystallization regime in the cores of small planetary bodies, encompassing conditions of the Moon and Ganymede.

2. Methods

2.1 In-situ CAESAR measurement under high pressure in a multi-anvil press

High-pressure experiments were carried out *in situ* in the DIA-type multi-anvil apparatus installed on the beamline PSICHE of SOLEIL, France. Starting materials with final compositions of 5, 10, 15, 20, 25 and 36.4 wt.% S, corresponding to 8.4, 16.3, 23.6, 30.4, 36.8 and 50 at.%, respectively, were obtained by a dry homogenization of mixtures of Fe and FeS powders in a mortar (without lubricants such as ethanol or acetone). The prepared mixtures were stored in a portable vacuum box to minimize the moisture absorption from air, and kept there until loading right before the experiment. For the high P-T experiments samples were placed in sapphire capsule capped with BN. The pressure assembly consisted of 26 mm WC cubes with 4 mm truncation and and pressure was measured *in situ* using the known equation of state of the MgO marker (Tange et al., 2009). Schematic view of the cell assembly is shown in Fig. S1. Trimmed octahedral pressure medium made from B+15 wt.% MgO (B85) was employed to minimize the absorption and diffraction of the sample environment (Xie et al., 2020). B85 gaskets were also used along the X-ray path instead of pyrophyllite gaskets. To remove the water absorbed during fabrication prior to the

experiment, MgO and ZrO_2 parts were baked at 1273 K for 1 hour; BN, graphite and B85 were kept at 393 K in a vacuum oven for more than 12 hours. Such a tailored cell assembly was proven to be fundamental to the collection of high-quality CEASAR data. High temperature was generated with graphite heaters and monitored with a $W_{97}Re_3$-$W_{75}Re_{25}$ thermocouple whose junction was set at the position symmetrical to the sample capsule with respect to the center of furnace. Status of cell assembly during the experiments was monitored by X-ray radiography.

Energy dispersive X-ray diffraction (EDD) measurements were carried out using a polychromatic X-ray (20-100 keV) focused to 25 μm vertically (FWHM) and collimated to 50 μm horizontally, much smaller than the sample dimension (0.8 mm diameter, 0.5 mm thick, Fig. S1). A Canberra SSD Ge detector associated with a XIA multi-channel analyzer with 2048 energy bins was used to acquire EDD covering energies up to 102.4 keV. The energies were calibrated using characteristic fluorescence X-ray lines of Mo, Sn, Ba, Sm and Au. The 2θ angle was calibrated from 5 to 25° using a 7 μm thick Au foil with a precision of ±0.003 °.

For all runs, the specimen was first brought to the target pressure and then heated to high temperature while collecting energy dispersive diffraction patterns at a fixed angle (8°). The melting of sample was identified from the disappearance of sharp diffraction peaks during temperature increase, and confirmed by the absence of any sharp peak during CAESAR acquisition. Once complete melting was achieved, CAESAR spectra were collected in order to obtain the structural and density information of the liquid. CAESAR collections were then repeated for increasing temperature, with 100-200 K steps. For each CAESAR scan, EDD were collected every 0.2° over a 2θ angle ranging from 2.5 to 24.1°, allowing acquisition of high quality raw data over a wide Q-range, up to 15 $A^{\circ -1}$ ($Q = 4\,\pi\,E\,\sin\theta\,/12.398$, where E is the energy of the X-rays in keV). All measured EDD are combined to form a two-dimensional array of intensities, Int(E, 2θ), with each Int value corresponding to a given E and 2θ index. Fig. 1a shows representative raw data collected on Fe-S in this study. CAESAR scan with good counting statistics were acquired with collection times of 5 sec for EDD between 2.5 and 10°, 10 sec for EDD between 10 and 20° and 20 sec for EDD between 20 and 24.1° (about 25 minutes in total). The horizontal gaps of the collimating slits were opened gradually during the angular scan from low angle to high angle to control the effective sample volume from which the diffracted X-rays was detected for discriminating sample signal from background scattering caused by surrounding materials. The raw CAESAR data were normalized by collection time and effective sample volume for further treatment. To improve counting statics, raw data were binned utilizing energy data following the method described in Wang et al. (2004) (Fig. 1b).

In view of the vertical focusing of the beam, 2D radiographies for alignment and sample observation purposes were recorded by scanning the press in front of the beam (scanning radiography) (Fig. 1c, d).

2.2 Analysis of the recovered samples

After experiments, the recovered cells were mounted in epoxy resin and polished along the plane parallel to the cylindrical axis of sample. The samples were polished sequentially

using diamond abrasive disks with 120, 30 µm and diamond paste with 1 µm grain size to obtain well-polished surface. Microstructure of recovered samples was analyzed by a field emission scanning electron microscope (SEM-FEG) (Zeiss Ultra55) at IMPMC, Sorbonne Université, France. The Fe and S contents and the concentrations of potential contaminants (B, N, C, Al, O, W, Re) were determined using electron probe microanalyses (EPMA) at Centre Camparis, Sorbonne Université, France using a Cameca SX-FIVE wavelength dispersive spectrometer (WDS) operating at 15 kV and 300 nA. Our recovered samples show fine dendritic textures of Fe and Fe-S; we therefore used a defocused beam of ~30 µm to average the compositions of the quenched liquid. Bulk chemical compositions were obtained by averaging 5-10 measurements. Table 1 reports the obtained average values and corresponding standard deviations.

2.3 Analysis of diffuse scattering signal

The scattering intensity curve, $I(Q)$, is constructed by merging the normalized EDD and removing the background (taking advantage of data collected at multiple 2θ angles over overlapping Q range (Fig. 2a)). The structure factor, S(Q), and the pair distribution function, $g(r)$, are calculated for the fixed Q range of 1-10 Å^{-1} for all cases. The structure factor, $S(Q)$, is first obtained after subtraction of incoherent scattering ($I_{inc}(Q)$). The distribution function, $F(r)$, and pair distribution function, $g(r)$, are obtained by the Fourier transformation of structure factor. Following the method detailed in (Morard et al., 2014, 2013), determination of density is based on the assumption that, due to the increasingly strong repulsive component in the interatomic potential, no atoms are located closer than minimal distance, r_{min}. Thus $F(r) = -4\pi r\rho$ for $r<r_{min}$. Density is, hence, extracted following the minimization of the oscillation in the short distance of the radial distribution function $g(r)$ (through a merit χ^2) (Eggert et al., 2002). Fig. 2b illustrates atomic densities calculated for different values of the minimal distance r_{min} at different temperatures. The value of χ^2 exhibits a well-defined minimum in most cases (r_0 in Fig. 2b), which gives the atomic density ρ_0 for which the relation $F(r) = -4\pi r\rho$ is best-satisfied (Fig. 2c). The position of r_{min} usually corresponds to the base of the first coordination sphere in $g(r)$. Noteworthy for the purpose of the present study, while absolute values of density depend upon r_{min}, density variation with temperature (dρ/dT) is effectively independent of r_{min} (Fig. 2b).

Various aspects of the data treatment enter into the assessment of the error bar on the absolute density value, including the selected Q range, the choice of the minimum distance of the first coordination sphere (r_0), along with physical phenomena neglected here such as the self-absorption from the sample (Morard et al., 2018, 2014). In the present data set, the estimated error is ±3 atoms/nm^3 for the atomic density, corresponding to ±220-270 kg/m^3 for the mass density of the Fe-S liquids (depending on composition).

2.4 Analysis at constant pressure

For analysis at constant pressure, experimental densities have been rescaled to 7 GPa following a Murnaghan formalism:

$$P - P_{ref} = \frac{K_{ref}}{K'}\left[\left(\frac{\rho}{\rho_{ref}}\right)^{K'} - 1\right] \quad (3)$$

where P is 7 GPa and ρ is the recalculated density at 7 GPa, P_{ref} and ρ_{ref} are our direct obtained experimental data at actual temperatures (Table 1), K_{ref} is the isothermal bulk modulus at P_{ref} and 1900 K, and K' is its pressure derivative. When not otherwise specified, we used the parameterization discussed in (Morard et al., 2018). Both K_{ref} and K' are functions of atomic S content (XS) in the liquid. The isothermal bulk modulus at ambient pressure and 1900 K of Fe-S alloys was obtained following (Chen et al., 2014) according to the relation $K_0 = K_{Fe}^{1-X_S} * K_S^{X_S}$, assuming an exponential dependence of the bulk modulus of liquid Fe-S alloys with S content, and with K_{Fe}=76 GPa and K_S=1.6 GPa. The pressure derivative of the bulk modulus (K') was obtained as $K' = K'_{Fe} + XS \cdot 3$, with K'_{Fe} = 6.5 (Morard et al., 2018). In this analysis the temperature dependence of bulk modulus was neglected (temperature range of interest < 500 K).

2.5 Thermodynamic models

To model the thermodynamic properties of liquid Fe-S (for brevity l-Fe-S) alloys we use the non-ideal solution model introduced in Morard et al. (2018) and Terasaki et al. (2019). The excess mixing volume is assumed to be pressure dependent and parameterized by an asymmetric Margules formulation. The two end-member phases of the binary solution model are l-Fe and l-FeS. The volume of the non-ideal solution is:

$$V(\chi_{FeS}, P, T) = (1 - \chi_{FeS})V_{Fe}(P,T) + \chi_{FeS}\, V_{FeS}(P,T) + \chi_{FeS}(1 - \chi_{FeS})\, V_{ex}(\chi_{FeS}, P) \quad (4)$$

where χ_{FeS} is the mol fraction of FeS, V_{Fe} and V_{FeS} are the molar volumes of pure l-Fe and l-FeS, and V_{ex} is the pressure and composition dependent excess mixing volume given by:

$$V_{ex}(\chi_{FeS}, P) = (\chi_{FeS}\, W_{Fe} + (1 - \chi_{FeS})\, W_{FeS})\, v(P), \quad (5)$$

where $v(P)$ is the pressure dependent contribution to the Margules parameters. Differently from Morard et al. (2018) and Terasaki et al. (2019) $v(P)$ is not parameterized by empirical expressions, but with a pseudospinodal equation of state (EOS) (Baonza et al., 1995):

$$v(P) = exp\left[\frac{1}{K'}\left(1 - \sqrt{1 + 2\,\frac{K'}{K_0}P}\right)\right] \quad (6)$$

The parameters W_{Fe}, W_{FeS}, K_0, and K' and the equation of state parameters of l-FeS are estimated from the data of this study as well as from density measurements of Morard et al. (2018) and acoustic velocities of Nasch et al. (1994) and Nishida et al. (2016).

Several recently published EOS of l-Fe (e.g. Komabayashi, 2014; Dorogokupets et al.,

2017; Wagle and Steinle-Neumann, 2019) are in relatively good agreement with respect to the prediction of density and acoustic velocities along isentropes at high pressure and temperature (Fig. S2a). Nonetheless, quantities derived from the associated thermodynamic potentials, such as the thermal expansivity and heat capacity that determine the adiabatic gradient in a convecting liquid core of a planet, show relatively large deviations (Fig. S2b, c). Therefore, the EOS of l-Fe affects not only the estimation of the temperature in the core but, together with dependent thermodynamic quantities, also determines the locations at which the core temperature drops below the liquidus. For this reason we assess here how the EOS of l-Fe of Komabayashi (2014), Dorogokupets et al. (2017) and Wagle and Steinle-Neumann (2019) affect the thermodynamic properties predicted by our l-Fe-S model. In the following we will refer to Model K, Model D and Model W to denote the thermodynamic models constructed on the basis of our datasets in the Fe-S system at high pressure and high temperature, which respectively make use of Komabayashi (2014), Dorogokupets et al. (2017) and Wagle and Steinle-Neumann (2019) for the EOS of l-Fe.

For the EOS of l-FeS we use a Vinet equation and the Anderson-Grüneisen formulation to describe the pressure dependence of the thermal expansivity (α) (e.g. Komabayashi, 2014). As we could not estimate α and the value of the Grüneiesen parameter at reference conditions ($\mathrm{P_{ref}} = 0.1\mathrm{MPa}$ and $\mathrm{T_{ref}} = 1650\,\mathrm{K}$) from our experimental data, we assume $\alpha = 11.8 \times 10^{-5}$ 1/K from Kaiura and Toguri (1979) and the isobaric heat capacity $C_P = 62.5$ J/K/mol (Chase, 1998). The Grüneisen parameter at reference conditions can then be computed from the estimated EOS parameters with the thermodynamic relation

$$\gamma = \frac{K_T V \alpha}{C_P - K_T V\, T\, \alpha^2}, \qquad (7)$$

where K_T is the isothermal bulk modulus and V is the molar volume. The EOS parameters and Margules coefficients estimated from the experimental data for the three different models, using different EOS of l-Fe, are given in Table S1.

The liquidus at sulfur concentration lower than that of the eutectic is parameterized following Buono and Walker (2011) (Eq. 7-9). This procedure neglects the small amount of S that can be dissolved in solid Fe and requires only the knowledge of the Gibbs energy of l-Fe and stable solid Fe phases together with the Margules parameters required to describe the non-ideal behavior of the Fe-S system. To obtain a thermodynamically consistent description we compute the liquidus according to the l-Fe EOS' of Komabayashi 2014 and Dorogokupets et al. 2017 together with the EOS' of the relevant solid phase provided by those authors (Model K and Model D).

Using those EOS' the Margules parameters are then estimated from experimental melting data. Here we make use of the experimental data at 1bar (e.g. Waldner and Pelton, 2005), 3 GPa (Brett and Bell, 1969), 6 GPa (Buono and Walker, 2011) and 10 GPa (Chen et al., 2008). With the knowledge of the Margules parameters and with the EOS' of l-Fe and solid Fe, the liquidus can then be computed at the required pressures. The estimated values of the Margules parameters are given in Table S2.

3. Results

The experimental data obtained here are summarized in Table 1 and described more in detail in the following subsections.

3.1 Recovered sample and chemical compositions

A representative cross-section of the recovered cell is shown in Fig. 3a. All of the samples (except FeS end member) exhibited homogeneous dendritic textures (Fig. 3c-g), characteristic of completely molten samples. The compositions of the recovered samples determined by EPMA marginally deviate from the starting compositions of the initial powder mixtures, showing a systematic increase by 0.7-1.8 at.% in the S content (Table 1). Potential causes include Fe exsolution from the liquid as an oxide, or simply a deviation from the expected 1:1 molar ratio in the FeS starting material. Possible chemical contamination by B, N, C, Al, O, W and Re was carefully checked. No Al contamination from sapphire capsule was ever observed for any of the samples. Some of the runs on samples with high melting temperature (low S content) showed B and/or N contamination from the BN cap (Fig. S1). For runs for which we experienced a thermocouple failure, often in relation to sample leaking as observed by *in-situ* radiography, W and Re were often detected in the analysis of the recovered experiments. In this paper we only present analysis of data from samples that show no contamination.

3.2 Structure of liquid Fe-S alloys

Fig. 4 shows examples of the structure factors, $S(Q)$, and derived pair distribution functions, $g(r)$, of liquid Fe-S alloys at high pressure and high temperature. The shape of $g(r)$ is characterized by two peaks located at approximately 2.35-2.55 Å and 4.7 Å corresponding to the distance of the first and second coordination spheres (CS), respectively. The sharpness of these features decrease with increasing S content, and a third local maximum around 6.5 Å is visible only for alloys with 31.7 at.% or less (the $g(r)$ become less structured with increasing S content). The first peak position of $g(r)$, r_1, was observed to monotonically decrease with increasing S content (Table 1, Fig. S3), which agrees quite well with previous report on liquid Fe-S (Shibazaki and Kono, 2018). The progressive reduction of r_1 with S content can be rationalized by considering that the partial pair distribution functions systematically show Fe-S bonds (~ 2.2 Å) shorter than Fe-Fe bonds (~ 2.5 Å) (Morard et al., 2018). As mentioned, with increasing S content, the first and second peaks of $S(Q)$ become broader, and oscillation of the $g(r)$ becomes less pronounced (Fig. 4a, b). The second oscillation in the $g(r)$ seems to vanish at 51.2 at.% S.

For a fixed S concentration, the $S(Q)$ and $g(r)$ do not show major changes with temperature (Fig. 4c, d). On a qualitative ground, this indicates that the temperature does not significantly modify the structure of liquid Fe-S alloys over the investigated range, other than for the increasing bond length.

3.3 *Density and thermal expansion of liquid Fe-S alloy*

All densities determined in this study are reported in Table 1. In selected cases we repeated the measurements in more than one heating cycle (e.g. MA40), showing a remarkable consistency. Also, using the nominal starting composition, the densities derived from data obtained in an experiment with leakage of the molten sample (MA58), yielded values compatible with those from more reliable data (MA66) obtained in the absence of leakage. The systematic consistency of our results proves the validity of our data collection strategy and data treatment (e.g. range of Q, value of r_{min}). As such, pressure and temperature derivatives of density, as its variation with S content, are highly reliable (see also discussion in section 2.3).

Previous density data for liquid Fe-S alloys under high pressure includes results obtained by ex-situ sink-float (Balog et al., 2003; Nishida et al., 2008), in situ X-ray absorption (Chen et al., 2014; Nishida et al., 2011; Sanloup et al., 2000; Terasaki et al., 2019) and X-ray diffraction (Morard et al., 2018) methods. The reported values display large discrepancies, however, even when the same method was used. Overall, the derived thermodynamic model from this study is compatible with densities determined here, although the agreement is less good for alloys with moderate S content (18 at.% or below) (Fig. S4). Of the parameters entering into the density modeling, the weight of thermal expansion is small, in particular when compared to that of compressibility. The most direct way to assess thermal expansion is thus considering the density evolution with increasing temperatures as from equation (2). Density determination for liquid Fe-S alloys with different S content at 7 GPa over an extended temperature range in this study is summarized and compared with thermodynamic Model W (Fig. 5).

For each composition, density decreases approximately linearly with increasing temperature (see section 2.4, equation 2). Aside from the already mentioned difference in the absolute density values noticeable for the alloys with low-S content, the thermodynamic model, is in good agreement with the experimentally observations for all the liquid alloys. Within the scatter of the data, the temperature derivative of the density doesn't show a clear compositional dependence (Fig. S5). As such, we cannot discriminate between results from thermodynamic Model K, Model D and Model W (see section 2.5), which equally well account for the measurements. Similarly, the experimentally derived thermal expansion shows values scattered in the range 6-13×10^{-5} K^{-1} (Fig. 6), in overall agreement with the results of the three thermodynamic models, but higher than previous models (Morard et al., 2018; Terasaki et al., 2019), in particular for S-rich samples.

To a large extent, the variability in our thermodynamic models reflects the spread in the values of thermal expansion of liquid Fe at ambient pressure (e.g. Williams, 2019). The density of model W has the strongest dependence on temperature for liquid Fe at ambient pressure (Fig. S6), with a thermal expansivity (14.1×10^{-5} K^{-1} at 1900 K) in agreement with Assael et al., (2006), while the thermal expansivity of Model K and Model D are smaller (respectively 9.0 and 9.3×10-5 K-1 at 1900 K) and in good agreement with the values of Hixon (1990).

4. Discussion

4.1 *Slope of adiabat and core solidification*

As already mentioned, the dynamic of crystallization in liquid planetary cores can be evaluated by comparing the adiabatic temperature gradient (slope of the isentrope) and the slope of the liquidus, which, in turn, depend upon pressure and chemical composition of the studied cores.

The adiabatic gradient in a liquid core is calculated according to the relation (1) for the three thermodynamic models, with densities benchmarked against measurements. Specific to liquid FeS, we use C_P=62.5 J/K/mol from NIST at 1bar (Chase, 1998). Assuming a different value of 40 J/K/mol (Kaiura and Toguri, 1979) does not significanlty change the resulting densities, slightly worsening the agreement with respect to experimental results, and making the isentropes moderately steeper.

The slopes of the adiabats obtained at selected pressures as a function of sulfur content are illustrated in Fig. 7 and compared with the slope of the liquidi (see section 2.5). It should be noted that the S concentrations considered here are always below the eutectic, where denser almost pure solid Fe is the phase crystallizing from the Fe-S liquid.

Observation of our different models highlights the dispersion depending on the chosen equation of state for liquid Fe. Nonetheless, on a qualitative ground, we note that the compositional domain for which the slope of the liquidus is larger than the adiabatic gradient is limited to a narrow portion on the Fe-rich side for all the pressure considered here (0-10 GPa).

On one hand, due to decreasing thermal expansion and increasing density, the slope of the adiabat decreases with increasing pressure. At the conditions of Moon's (5.2 GPa) and Ganymede's (10 GPa) center, the adiabat only moderately depends on S content, as a consequence of the density reduction that overcomes the effect on C_P (Fig. S2). On the other hand, the slope of the liquidus shows a significant decrease with increasing S content. The net effect is that, for Fe-S cores of small to middle size bodies (0 to 10 GPa) top-down solidification is the most likely scenario for wide range of S contents. Only for S-poor cases bottom-up inner grow is possible, like in the case of the Earth.

However, uncertainties on the relevant thermodynamic quantities resulting from the liquid equation of state of Fe and thermoelastic properties of l-Fe-S alloys do not allow derivation of quantitative conclusions on the crystallization regime of planetary cores as a function of composition. For instance, assuming Model W, crystallization will always occurs top-down because the adiabatic gradient is always steeper than the liquidus over the pressure range considered here (Fig. 7c). Conversely, assuming the Model K or Model D, adiabatic gradients and the slope of the liquidus cross each other at S concentrations that increase with increasing pressure (Fig. 7a-b).

Paleomagnetic studies of chondritic and small-body achondritic meteorites have revealed a large diversity of magnetic field records (e.g. Weiss et al., 2020). Convection in the cores of differentiated, or at least partially differentiated asteroids, planetesimals, or meteoritic parent bodies in general, may have generated dynamo magnetic fields responsible for the magnetization of the overlying silicate rocks (Gattacceca et al., 2016; Weiss et al., 2008).

While heating from short-living radioactive isotopes is generally considered to be responsible for the early dynamo, core solidification is expected to produce compositionally-driven dynamo activity,at a later stage with the timing dependent on the S concentration of the core and the radius of the body (Neufeld et al., 2019). At low pressures within an asteroid, isentropes are steeper than liquidus, irrespective of the S content for both models (b) and (c), while the adiabatic gradient is below the slope of the liquidus for alloys up to 4 at.% S according to model (a) (Fig. 7). Top-down crystallization thus remains the most likely scenario, but bottom-up cannot be firmly excluded for very low S content. Generation of a magnetic field during top-down core crystallization could therefore be relatively common in differentiated planetesimals in the early age of the solar system.

Lunar core formation models based on metal/silicate partitioning of siderophile elements support a Moon core containing up to 10 at.% S (Rai and Van Westrenen, 2014; Stenstra et al., 2017). Assuming a fully molten core in the binary Fe-S system, geodetic constraints argue for a larger amount of S, between 16 and 40 at.%, depending on the core radius (Morard et al., 2018). The presence of a solid inner core, made of pure Fe, brings estimates down below 10 at.% (Antonangeli et al., 2015). At about 5 GPa, the pressure at the center of the Moon, the adiabatic gradient and slope of the liquidus might cross at about 3 or about 9 at.% S, depending on the EOS of l-Fe (Fig. 7). Accordingly, if the Moon does not have a solid inner core today, it will have most likely crystallized top down. On the contrary, if the existence of a solid inner core, as suggested by some seismological studies (Weber et al., 2011), is confirmed, its crystallization regime remains uncertain. That an initial bottom-up scenario, may have evolved into a top-down regime due to the progressive enrichment in S of the liquid portion of the core upon crystallization of pure Fe, as hypothesized to explain the early, now extinct Moon magnetic field (e.g. Laneuville et al., 2014), remains an appealing possibility.

Compositional convection is considered a fundamental element to explain Ganymede's present-day dynamo (e.g. Rückriemen et al., 2018). In turn, compositional convection is strongly linked to core differentiation and solidification. Recently, significant efforts have been dedicated to model dynamics and magnetic field generation under the notion that crystallization occurs at the top of Ganymede's core (Breuer et al., 2015; Rückriemen et al., 2018, 2015). Our results suggest this is not necessarily the case, and that crystallization might proceed bottom-up for S content below about 10 at%. Indeed, the solution space for bottom-up crystallization might increase up to 7-11 at.% S, again depending on assumed liquid Fe equation of state, at 10 GPa, central pressure of Ganymede's core (Fig. 7). As both bottom-up and top-down crystallization can power a core dynamo and generate a magnetic field, although via different mechanisms, independent constraints on S abundance in Ganymede's core are needed to assess its thermo-chemical history.

5. Conclusions

We performed a structural investigation of liquid Fe-S alloys at high pressure over a large temperature range and compositional domain. Derived densities have been used to refine thermodynamic models of the thermo-elastic properties of the liquid Fe-FeS system

and, in particular, to put constraints on modeled thermal expansion at conditions directly relevant for the core of small planetary bodies.

Comparison of the adiabatic gradient with the slope of the liquidus, is used to discuss the crystallization regime of the core of planetary bodies in the range 0-10 GPa as a function of sulfur concentration, with specific emphasis to the case of asteroids and small planetesimals, the Moon, and Ganymede. Our results show that the compositional domain for which the slope of the liquidus is larger than the adiabatic gradient is limited to a narrow portion on the Fe-rich side for all the pressures considered here, implying that top-down crystallization is likely a widespread phenomenon. However, bottom-up crystallization is still possible for S-poor cases, and the compositional domain for which an inner core would grow bottom-up increases with increasing pressure. On the basis of our experimental results and thermodynamic modeling, a bottom-up scenario cannot be excluded to occur at the pressures of the core of the Moon and, all the more so, of Ganymede. Improved constraints on the liquid Fe EoS are needed to better discriminate possible scenarios.

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750 Table 1 Chemical composition of quenched samples, pressure-temperature conditions, and
751 measured densities and position of the first sharp peak *r*1. The density scaled to a constant
752 pressure of 7 GPa is also shown.

Run	S content (at.%)	Pressure (GPa)	Temperature (K)	Density (kg/m3)	Density corrected to 7 GPa (kg/m3)	*r*1 (Å)
MA84	51.2±0.4	6.5	1700	4486	4521	2.374
		6.8	1800	4427	4440	2.375
		6.8	1900	4332	4345	2.376
		6.5	2000	4286	4319	2.379
		5.7	2100	4222	4312	2.378
MA40	38.3±0.5	6.5	1280	5619	5659	2.440
		6.4	1400	5526	5578	2.442
		6.4	1500	5511	5562	2.442
		6.7	1320	5619	5642	2.446
		6.6	1310	5619	5654	2.446
		6.9	1500	5573	5577	2.440
		6.0	1600	5464	5545	2.446
		6.1	1790	5348	5420	2.439
MA44	31.7±0.9	7.1	1415	5944	5938	2.450
		6.9	1500	5880	5889	2.447
		6.8	1610	5800	5817	2.451
		7.0	1710	5720	5721	2.454
		6.9	1800	5656	5666	2.454
MA47	25.1±0.3	7.5	1600	6211	6171	2.485
		7.5	1705	6160	6119	2.485
		7.8	1790	6137	6077	2.489
		7.9	1900	6084	6018	2.486
		8.4	2050	6059	5956	2.484
		8.3	2020	6043	5951	2.486
MA82	18.1±0.4	5.6	1500	6308	6416	2.502
		5.3	1600	6291	6423	2.512
		5.4	1700	6188	6310	2.515
		5.2	1800	6171	6309	2.520
MA66	9.1±0.5	6.8	1850	6746	6759	2.541
		7.1	1950	6711	6704	2.553
		7.4	2050	6647	6621	2.555
		7.7	2155	6569	6524	2.554
		8.1	2250	6520	6451	2.554
MA58	8.4*	5.5	1825	6717	6880	2.537
		5.4	1960	6599	6766	2.541
		4.7	2075	6487	6703	2.554

753
754 § The uncertainty of ±3 atoms/nm^3 estimated for the used protocol in data analysis (Morard

et al., 2013), reflects into an error of 220, 230, 240, 250, 260 and 270 kg/m^3 for 51.2, 38.3, 31.7, 25.1, 18.1 and 9.1 at.% S content, respectively

* Starting composition; analysis of the recovered sample was not possible due to the leak of the sample at further high temperature.

FIGURES

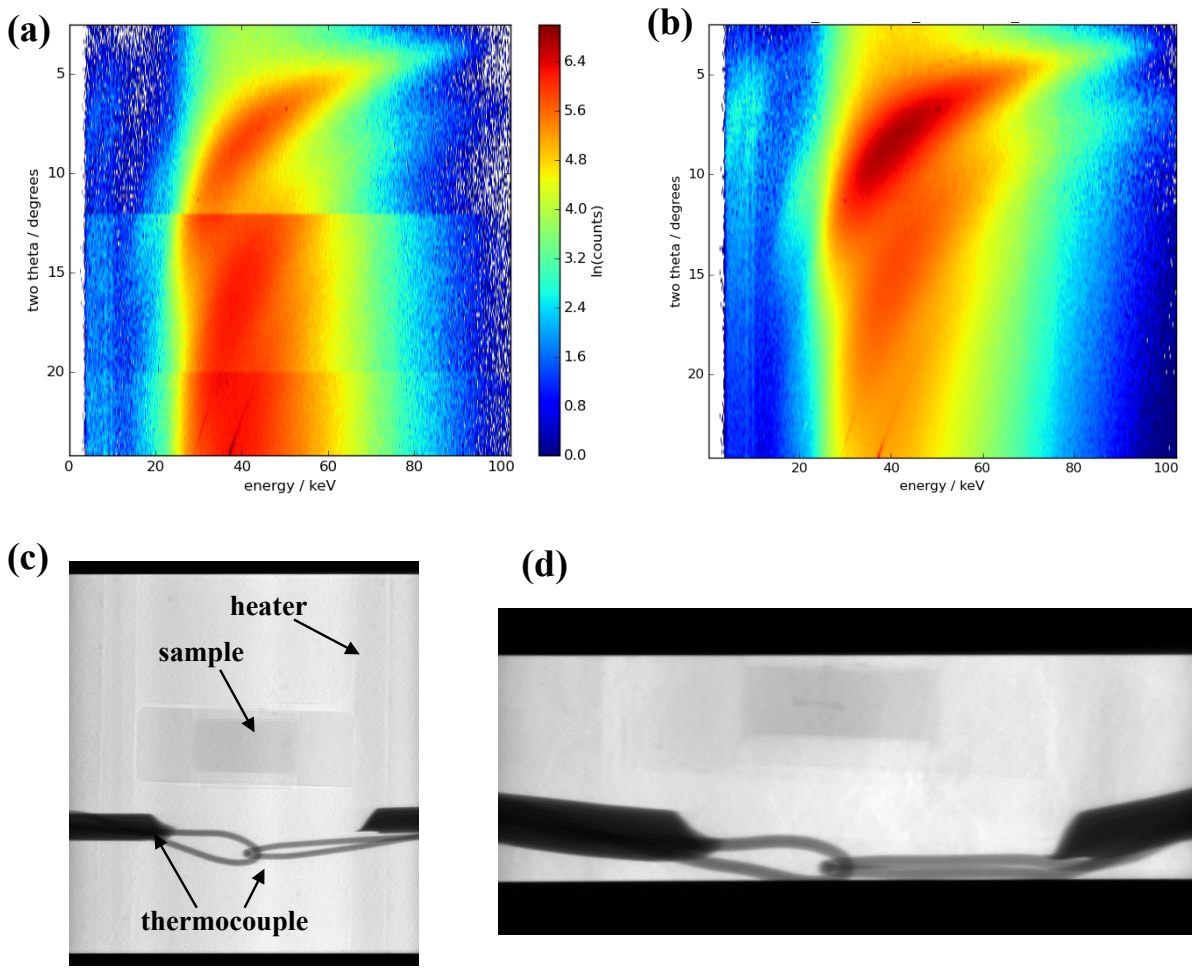


Fig. 1 (a, b) Two-dimensional CAESAR plot of diffraction intensities collected on liquid FeS (run MA84). Raw 2D data (EDD spectra as a function of angle) (a) and data after normalization for acquisition time and binning (b). The discontinuities visible in (a) for 2θ=12°and 20° are due to the change of acquisition time. (c, d) Images of cell assemblies obtained by X-ray radiography before compression (a) and at 6.3 GPa (c) after reaching to the target load.

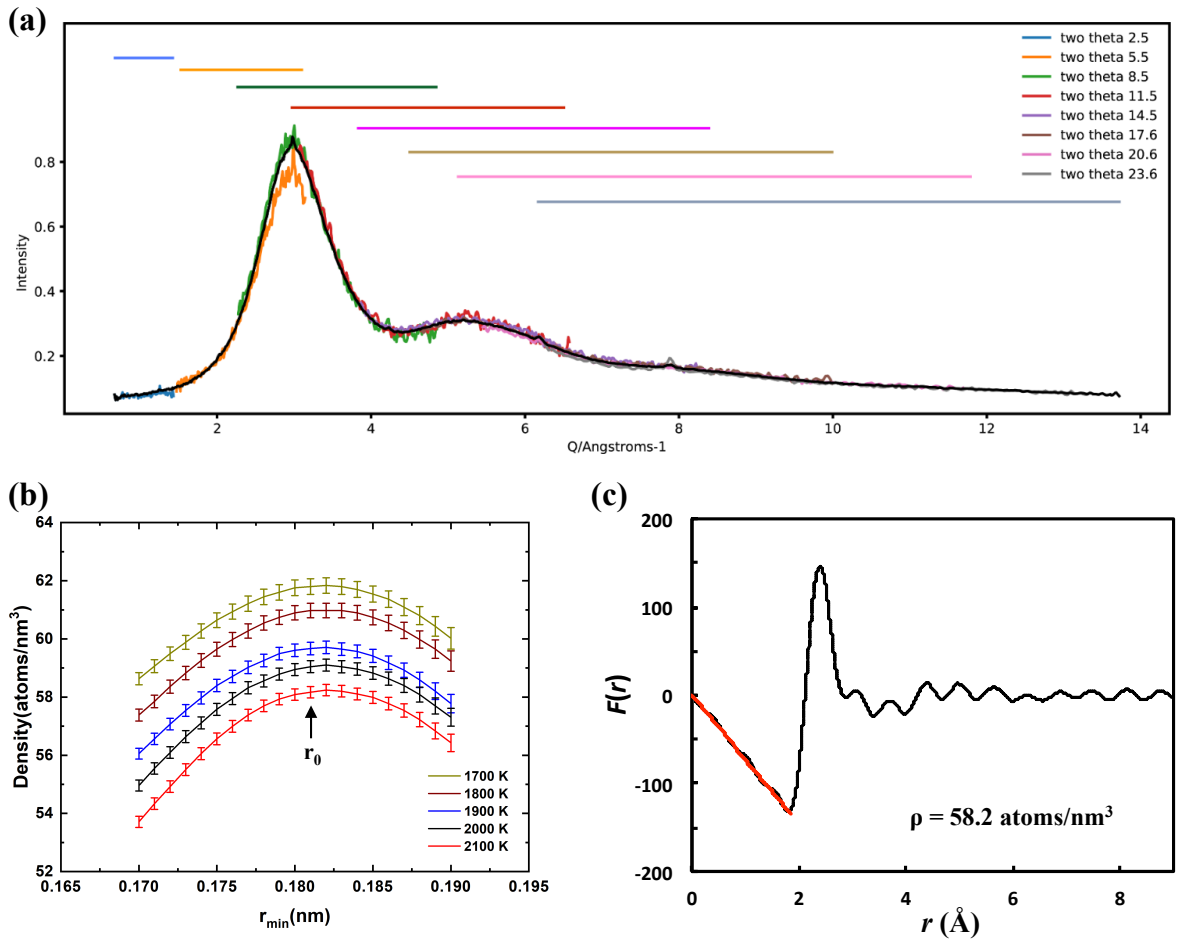
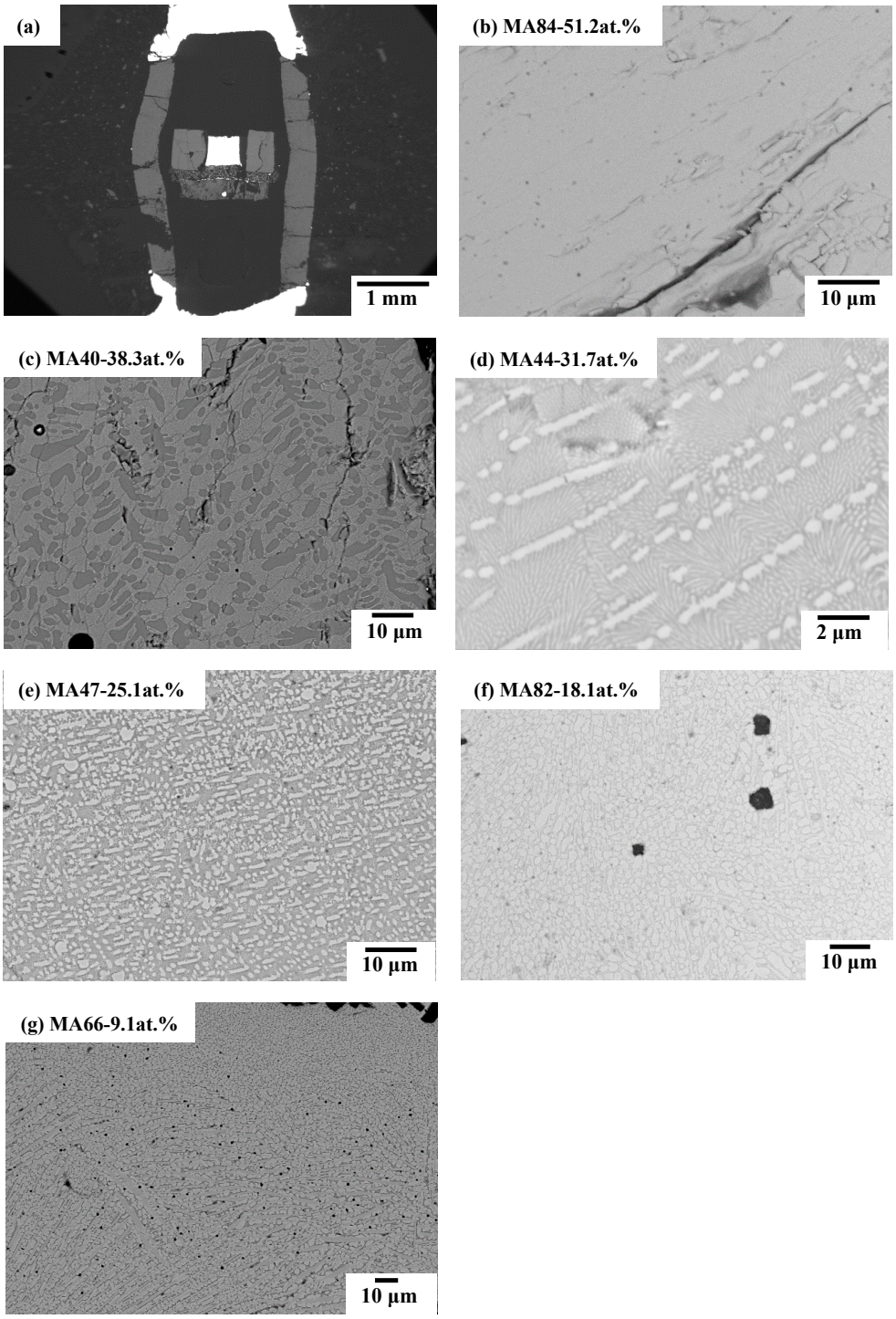


Fig. 2 Examples of raw CAESAR data treatment process. (a) Normalized scattering intensity plotted as a function of Q. Color-coded horizontal bars and overlapping I(Q) patterns illustrate the Q range covered and *I(Q)* derived from the EDX spectra at corresponding 2θ. (b) Atomic densities calculated for different values of the minimal distance r_{min}. Error bars indicate the value of merit χ^2 for each r_{min} (see Morard et al. (2013) for more details). Arrow points out the local minimum of χ^2, providing atomic density and r_0 for this data set. We stress that the temperature dependence of density is effectively independent of r_{min}. (c) Comparison of the distribution function *F(r)* with the function $-4\pi r\rho_0$ (dashed lines) calculated using density obtained by χ^2 minimization (d) (2100 K).



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795 **Fig. 3** Backscattered electron image of recovered samples. (a) An overall view of the
796 recovered cell after run MA47. (b-g) Microstructure of the recovered samples of the indicated
797 runs. In (c-g) brighter areas correspond to Fe-rich portions of the sample, while darker areas
798 correspond to S-rich portions.
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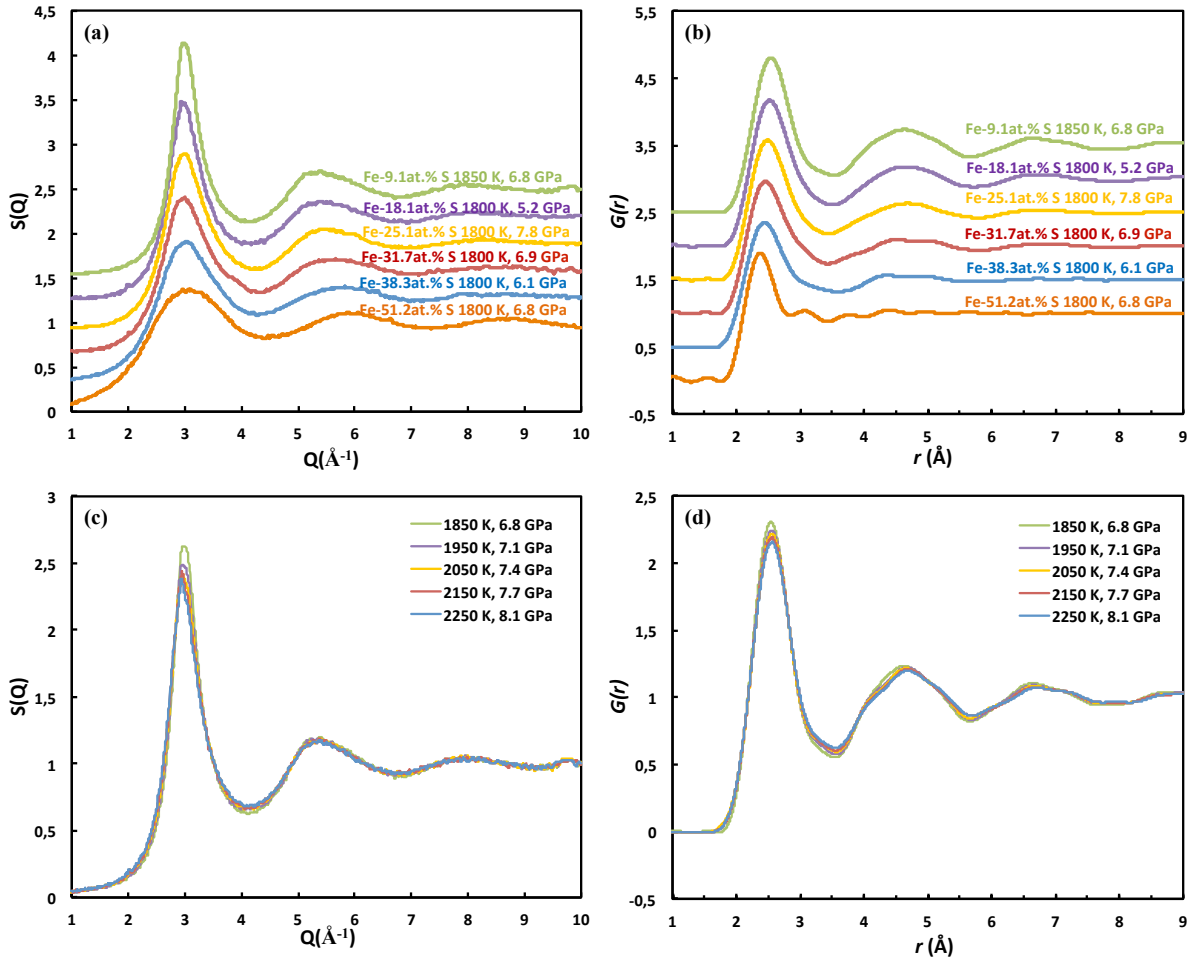


Fig. 4 Examples of structure factor, *S(Q)*, and pair distribution function, *g(r)*, of liquids Fe-S alloys at high pressure and high temperature. (a) *S(Q)* and (b) *g(r)* as a function of S content. For clarity, *S(Q)* and *g(r)* are shown with a vertical offset of 0.3 and 0.5 in (a) and (b), respectively. Oscillations at ~ 4.4 and 5.0 Å visible in (b) for S concentration larger than 31.7at.% are not real features and arise from spurious signal due to the limited Q range in the corresponding structure factors. (c) *S(Q)* and (d) *G(r)* of Fe-9.1at.%S as a function of temperature (MA66).

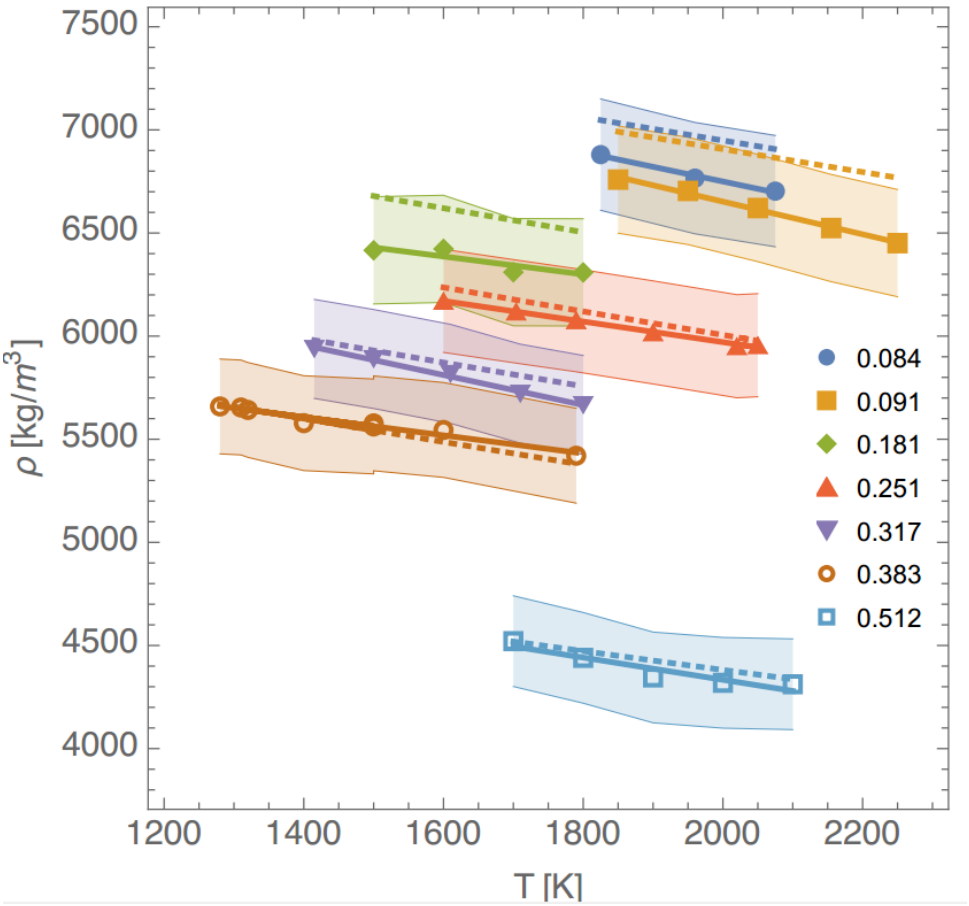


Fig. 5 Temperature dependence of the density of liquid Fe-S alloys at constant pressure (7 GPa). Linear fit of our data (solid line) and temperature derivative according to the thermodynamic Model W (dashed lines). Legend: atomic S concentration. Hatched areas indicate uncertainties on absolute density values (±3 atoms/nm^3, see section 2.3). Reported liquid compositions are those determined by chemical analysis of quenched and recovered samples but for the experiment on Fe-$S_{0.084}$ for which we considered the composition of the starting material.

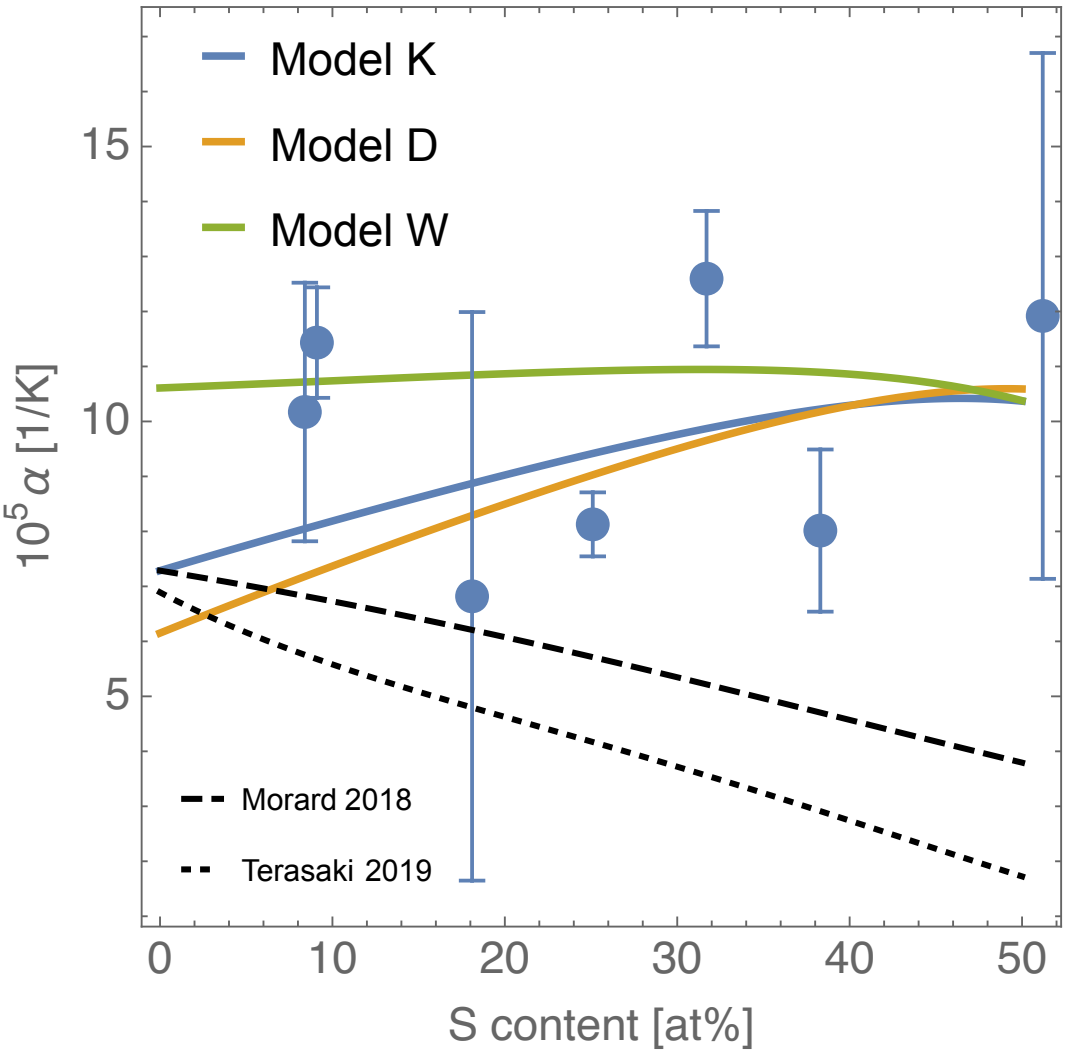


Fig. 6 Thermal expansivity as function of sulfur content at 7 GPa and 2000 K. Dots are estimations from experiments while colored lines are outcomes of thermodynamic models - Model K (blue), Model D (orange), and Model W) (green). Please refer to section 2.5 for details on the thermodynamic models and Table S1 and S2 for the used parameters. Black lines are results from literature thermodynamic models (dotted - Terasaki et al., 2019; dashed - Morard et al., 2018).

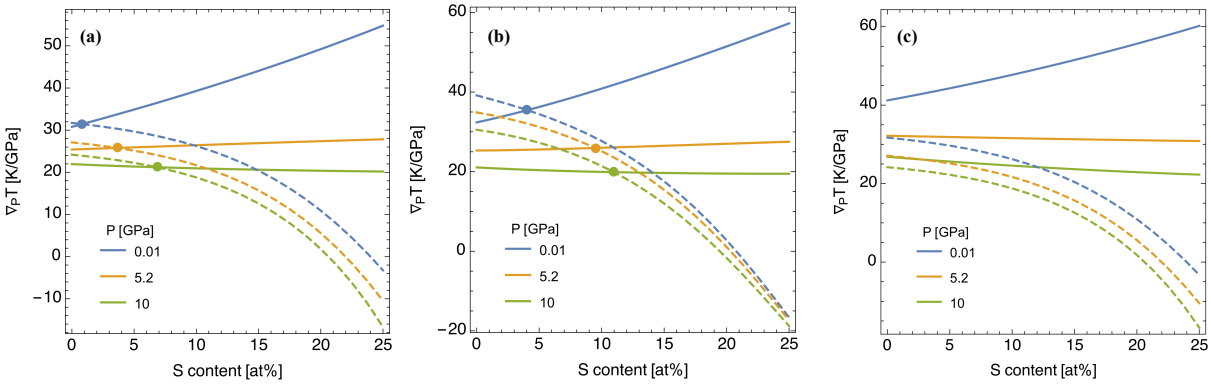


Fig. 7 Adiabatic gradient (solid lines) and slope of the liquidus (dashed lines) as a function of the S content at 0.01 GPa, 5.2 GPa, and 10 GPa, corresponding to core pressure of planetesimals, Moon and Ganymede, respectively, for the Model K (a), Model D (b), and Model W (2019) (c). Note that for the Model W the EOS' of solid Fe from Komabayashi (2014) has been used to compute the slope of the liquidi. For other differences in the models, please refer to section 2.5 and to Table S1 and S2 for the used parameters.

Supplementary material for:

Thermal expansion of liquid Fe-S alloy at high pressure

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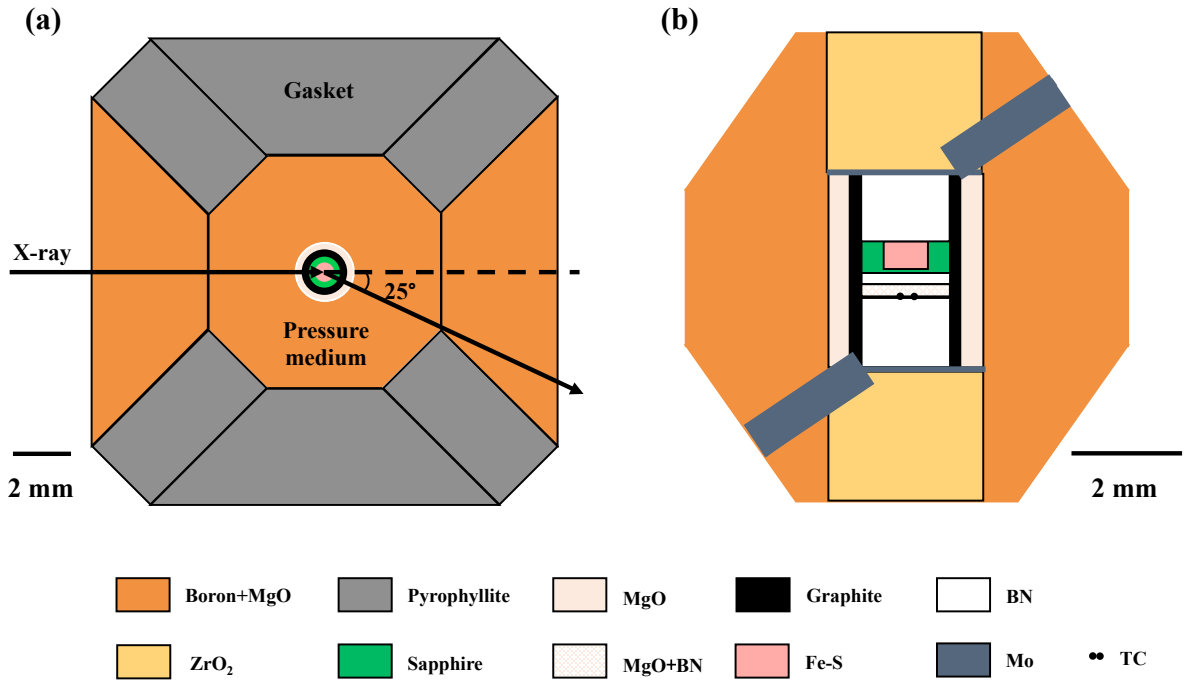


Fig. S1 Schematic illustration of the experimental design. (a) A top view of the diffraction geometry through the gasket and cell assembly. (b) A cross-section of cell assembly used in this study. Temperature was monitored with a $W_{97}Re3$-$W_{75}Re_{25}$ thermocouple (TC) whose junction was indicated by black dots.

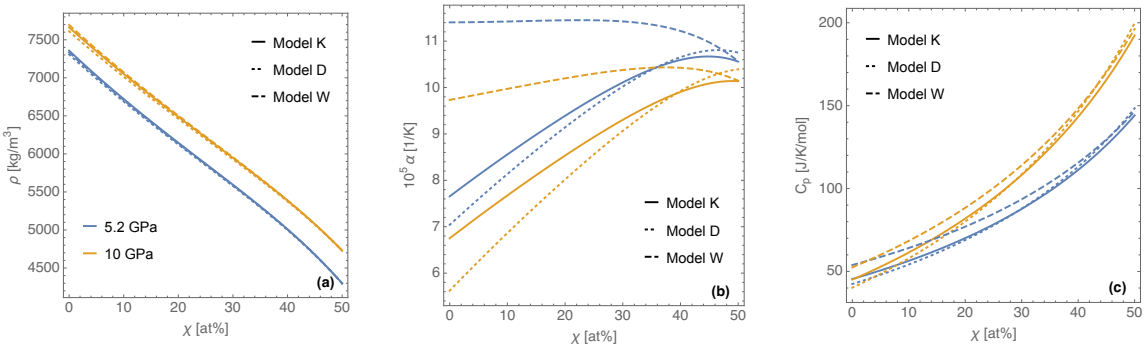
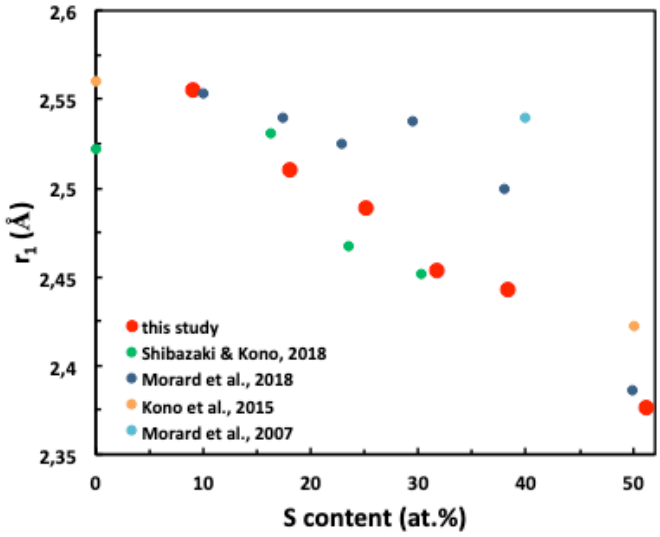


Fig.S2. Density (a), thermal expansivity (b), and iso-baric heat capacity (c) as a function of molar sulfur fraction X at 5.2 GPa and 10 GPa and 2000K for the thermodynamic Model K, Model D and Model W. Model K, Model D and Model W differ in the reference EOS for liquid iron (Model K uses l-Fe EOS of Komabayashi (2014), Model D uses l-Fe EOS of Dorogokupets et al. (2017), Model W uses l-Fe EOS of Wagle and Steinle-Neumann (2019)). Parameters of the thermodynamic models are in Table S1 and S2.



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907
908 **Fig.S3.** First peak position (r_1) as a function of the S content. Previous results from
909 Shibazaki et al. (2018) at 3-5 GPa, Morard et al. (2018) at 2-5 GPa, Kono et al. (2015) at 1-6
910 GPa, and Morard et al. (2007) at 3-17 GPa, were also shown for comparison. For clarity only
911 one value is plotted for each of the considered composition (r_1 for a given S content only
912 moderately depends on pressure and temperature over the P-T range covered by individual
913 studies).
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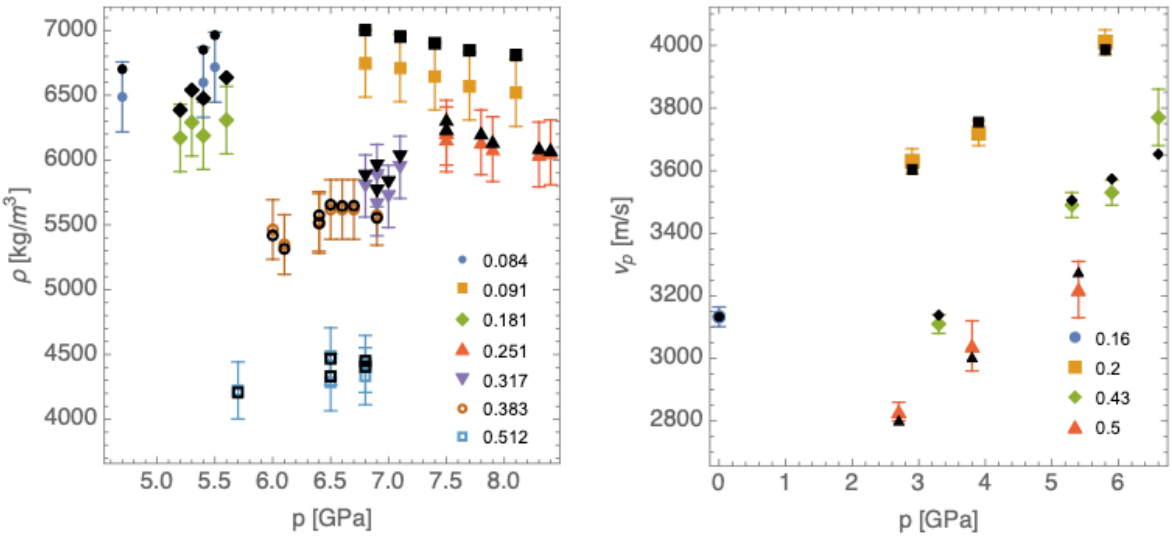


Fig. S4. Measured densities (this study) and acoustic velocities (Nasch et al. 1994 and Nishida et al. 2016) (colored symbols) at varying temperature and predicted values according to the thermodynamic Model W (black symbols) under the same condition. Different symbols corresponds to different S content. For details on the thermodynamic model, please refer to section 2.5 and to Table S1 and S2 for parameters.

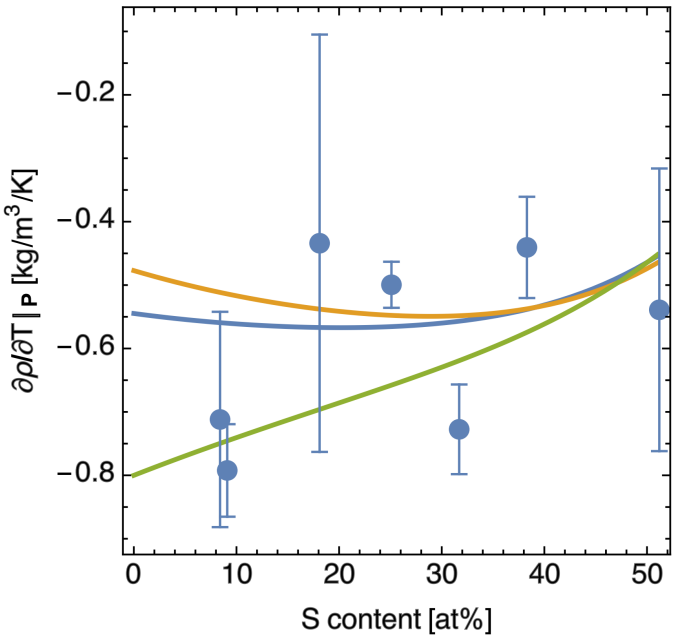


Fig. S5. Temperature derivatives at the constant pressure (7 GPa) of density as a function of S content. Dots are the slopes of the linear fit to the experimental data (see Fig. 6) with uncertainties at 2σ. Colored lines are outcomes of thermodynamic Model K (blue), Model D (orange) and Model W (green). Model K, Model D and Model W differ in the reference EOS for liquid iron (Model K uses l-Fe EOS of Komabayashi (2014), Model D uses l-Fe EOS of Dorogokupets et al. (2017), Model W uses l-Fe EOS of Wagle and Steinle-Neumann (2019)). Parameters of the thermodynamic models are in Table S1 and S2.

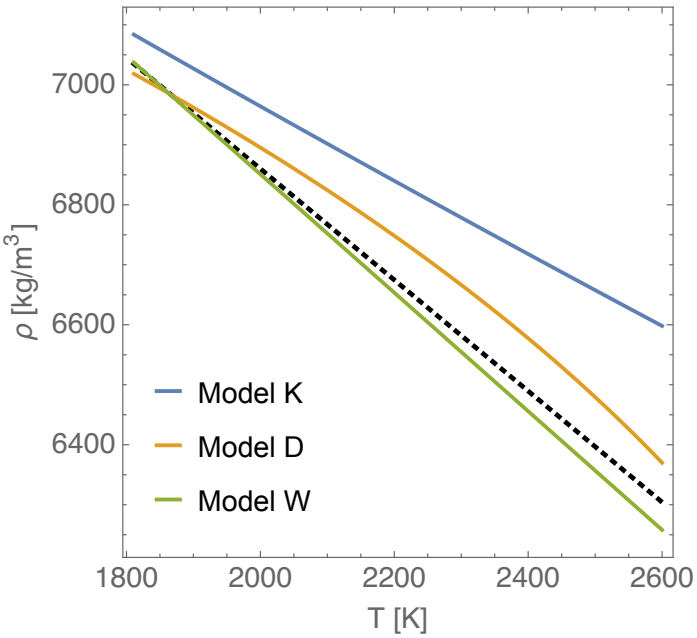


Fig. S6. Density as a function of temperature for liquid iron at ambient pressure. Colored lines are outcomes of our thermodynamic models (see section 2.5 and Table S1 and S3 for details). Dashed black line is a fit to data from Assael et al.. (2006), and preferred fit according to Williams (2009).

Table S1. EOS parameters for l-FeS and Margules parameters estimated from the experimental data of this study and from density measurements of Morad et al. (2018) and acoustic velocities from Nishida et al. (2016) using the EOS of l-Fe of Komabayashi (2014), Model K, Dorogokupets et al. (2017), Model D and Wagle and Steinle-Neumann (2019), Model W. $P_{ref} = 0.1$MPa, $T_{ref} = 1650$ K and for the l-FeS end-member $\kappa = 1.4$, $\alpha = 11.8 \times 10^{-5}$1/K, and $C_p = 62.5$ J/K/mol.

Thermodynamic model	l-Fe EOS	V [cm^3/mol]	K_T [GPa]	$K_T{}'$	γ	δ_T	W_{Fe}	W_{FeS}	B_0 [GPa]	B'
Model K	Komabayashi 2014	24.25	13.22	6.36	0.68	0.62	-9.627	-3.435	3.45	2.14
Model D	Dorogokupets 2017	24.26	13.12	6.38	0.68	0.52	-9.481	-3.528	3.44	2.10
Model W	Wagle 2019	24.26	13.07	6.44	0.68	0.61	-9.627	-3.393	3.44	2.23

Table S2. Margules parameters required to compute the Fe-S liquidus (see Eq. 9-11 Buono et al. 2011) for the EOS' of l-Fe of Komabayashi (2014) and Dorogokupets et al. (2017). For the Model W the EOS' of solid Fe from Komabayashi (2014) has been used to compute the slope of the liquidi.

Thermodynamic model	l-Fe EOS	$W_{H,Fe}$ [kJ/mol]	$W_{S,Fe}$ [kJ/mol/K]	$W_{V,Fe}$ [kJ/mol/GPa]	$W_{H,FeS}$ [kJ/mol]	$W_{H,FeS}$ [kJ/mol/K]	$W_{H,FeS}$ [kJ/mol/GPa]
Model K	Komabayashi 2014	53.70	0.029	-3.82	25.34	0.0	-2.95
Model D	Dorogokupets 2017	69.28	0.043	-2.78	27.3	0.0	-3.48